Astro 500

Techniques of Modern Observational Astrophysics

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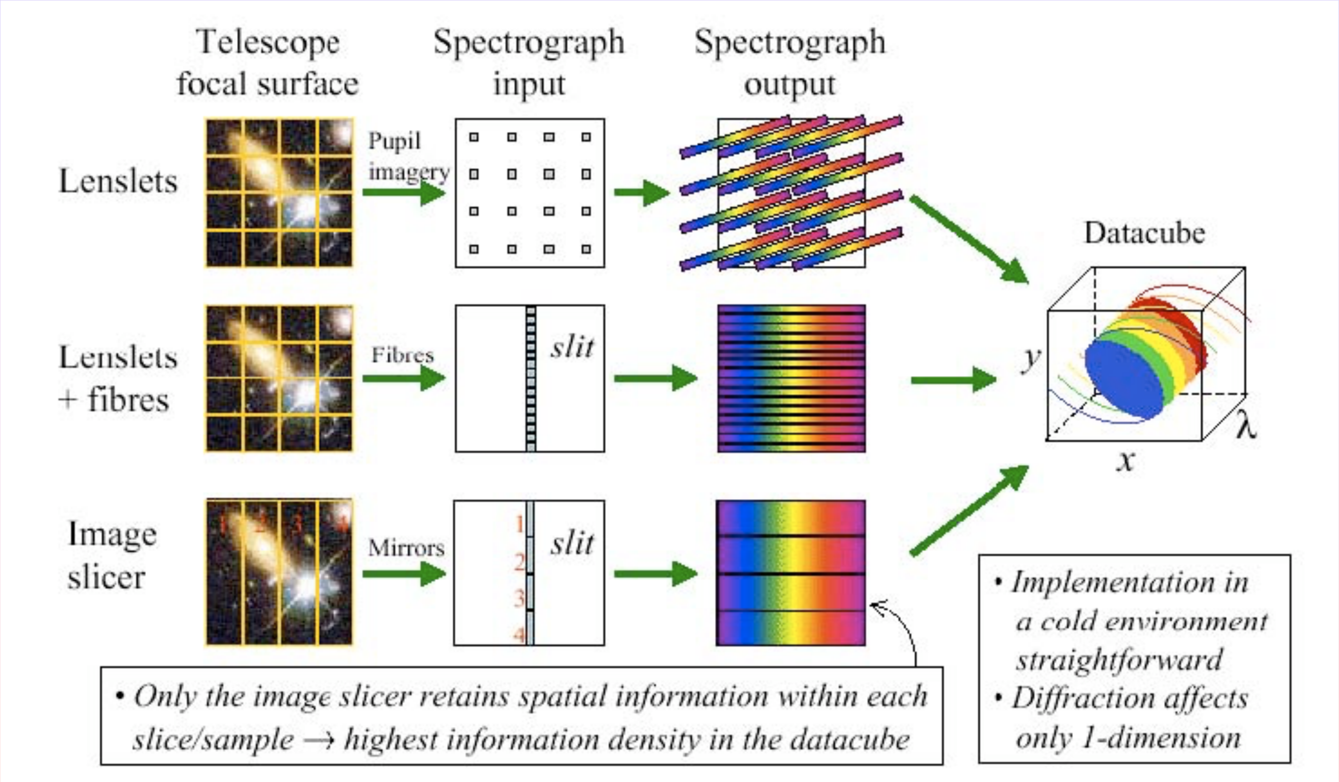
Approaches

Examples of available instruments

- ✓ Grating-dispersed spectrographs
  - ✓ basic spectrograph design
  - ✓ dispersive elements
  - ✓ Long-slit spectrographs
    - ✓ General Observing Considerations
  - ✓ Double spectrographs
  - ✓ Multi-objects spectrographs: slitlets vs fibers
  - ✓ Echelle spectrographs
  - ➢ 3D spectroscopy: coupling formats and methods
    - ✓ Fiber
    - o Fiber+lenslet
    - o Slicer
    - o Lenslet
    - o Filtered multi-slit
    - o 3D MOS
  - ➢ Current instruments
    - o summary of considerations
    - o sky subtraction

IFUs

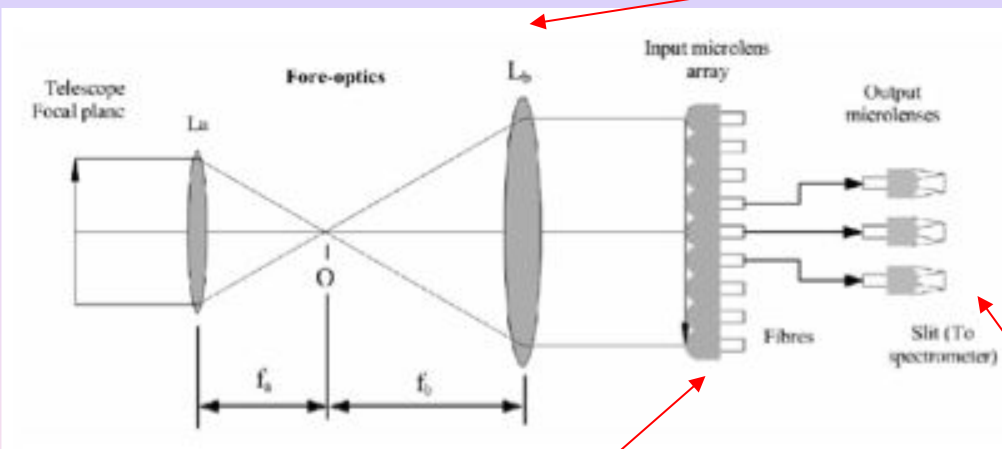


Telescope focal surface
Spectrograph input
Spectrograph output
Lenslets
Pupil imagery
Datacube
Lenslets + fibres
Fibres
slit
y
λ
x
Image slicer
Mirrors
1
2
3
4
slit
• Only the image slicer retains spatial information within each slice/sample → highest information density in the datacube
• Implementation in a cold environment straightforward
• Diffraction affects only 1-dimension

Grating-dispersed spectrographs

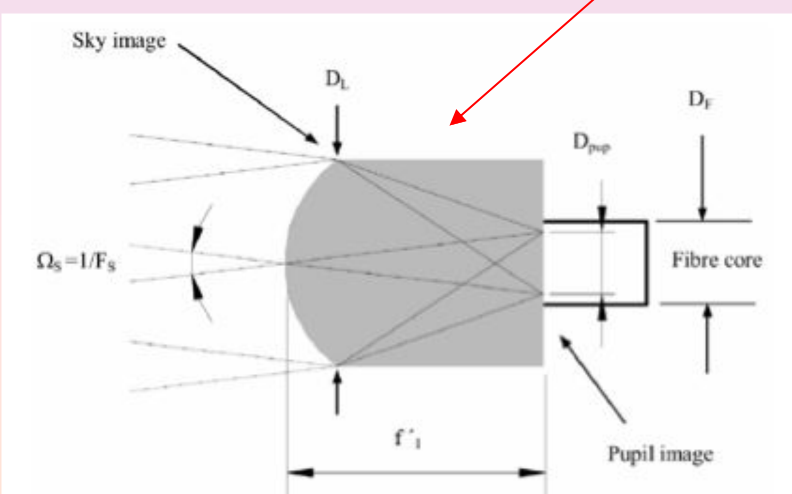
fibers + lenslet feeds

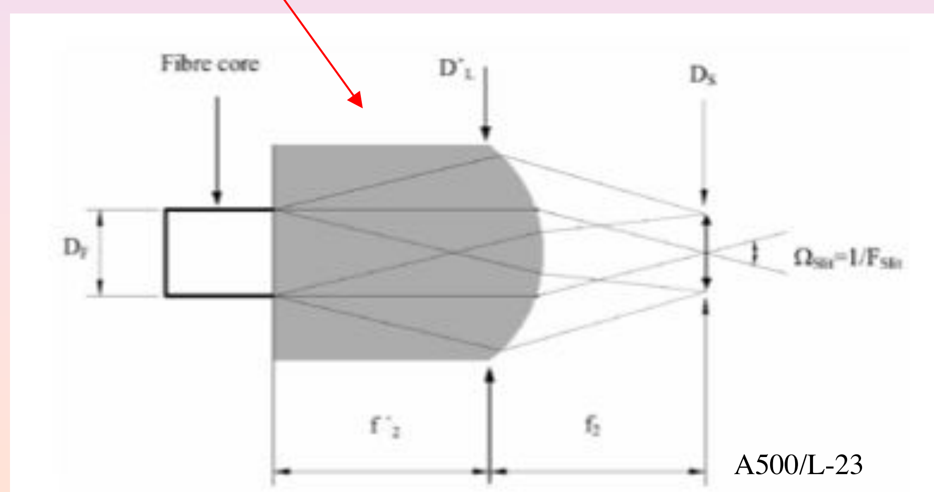
- concept:



- Focal exander matches to scale of lenslet array
- Micro-lens forms pupil image on fiber
- Pupil image is smaller; angles are larger (AΩ again)
- Option to reform slit-image with output micro-lens

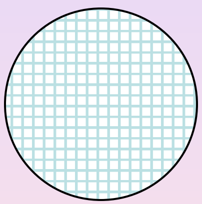
Ren & Allington-Smith '02

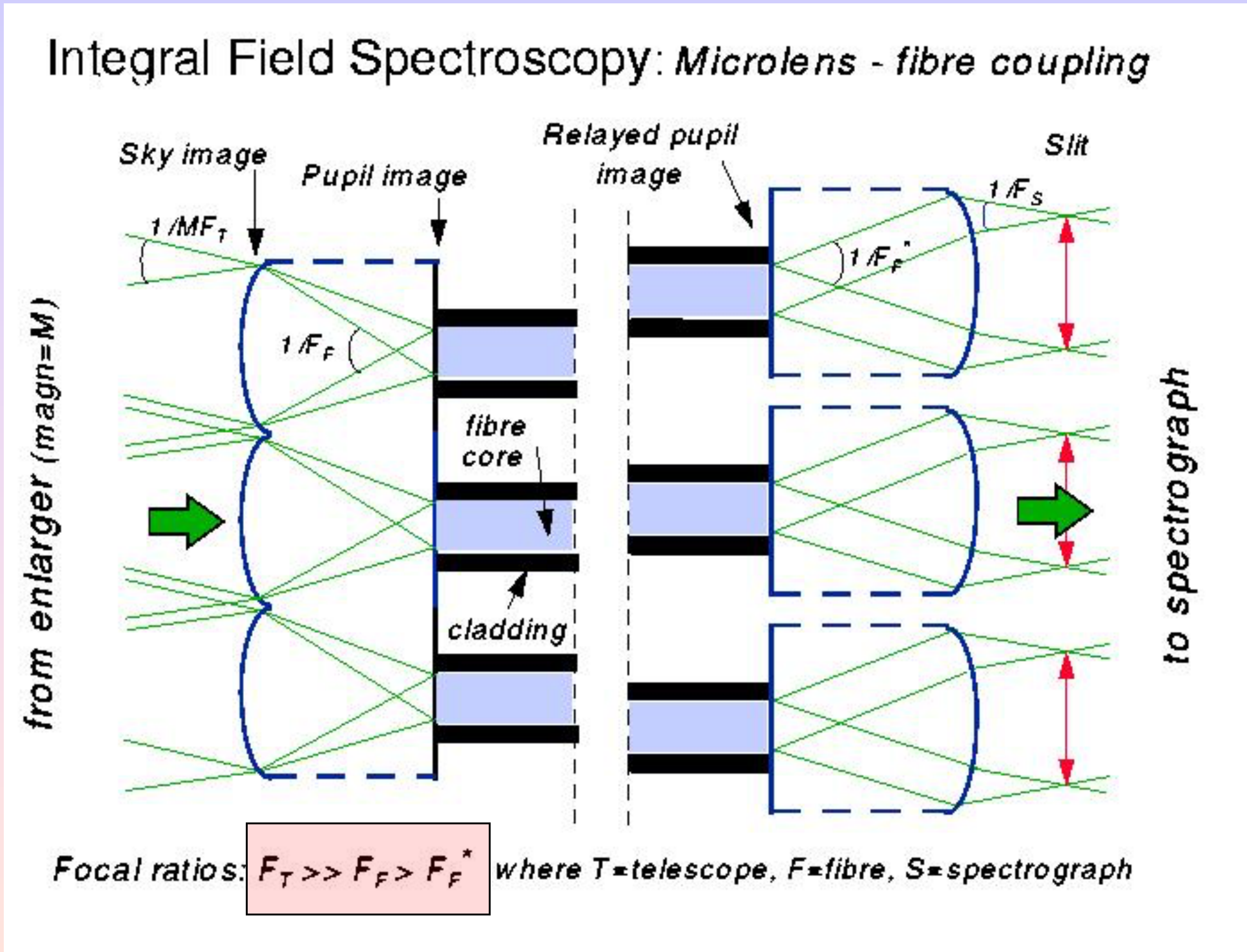




Grating-dispersed spectrographs

fibers + lenslet feeds





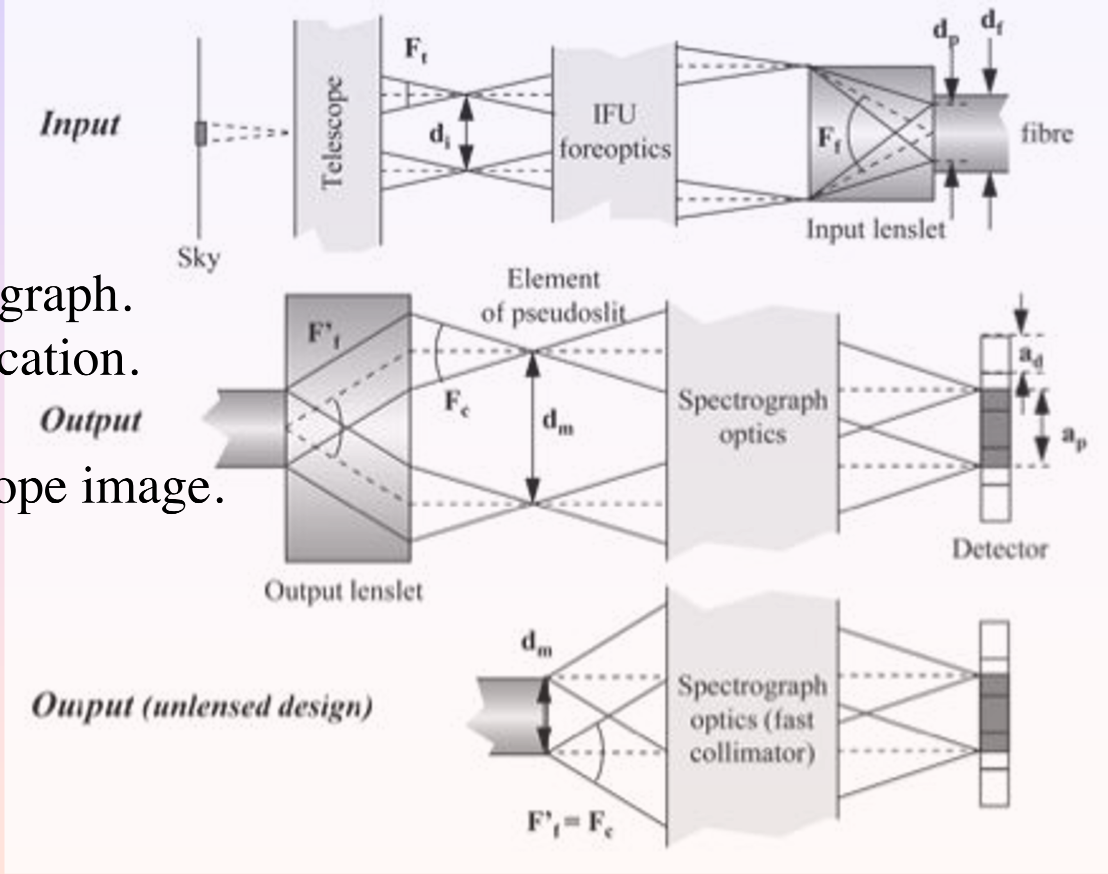
Allington-Smith & Content '98

Grating-dispersed spectrographs
fibers + lenslet feeds

Allington-Smith & Content '98

Note: don't have to use lenslets at output end:

- Higher input f-ratio to spectrograph.
- Less possibility for demagnification.
- Spectrograph images pupil.
- Ray-bundle varies with telescope image.



Grating-dispersed spectrographs
fibers + lenslet feeds

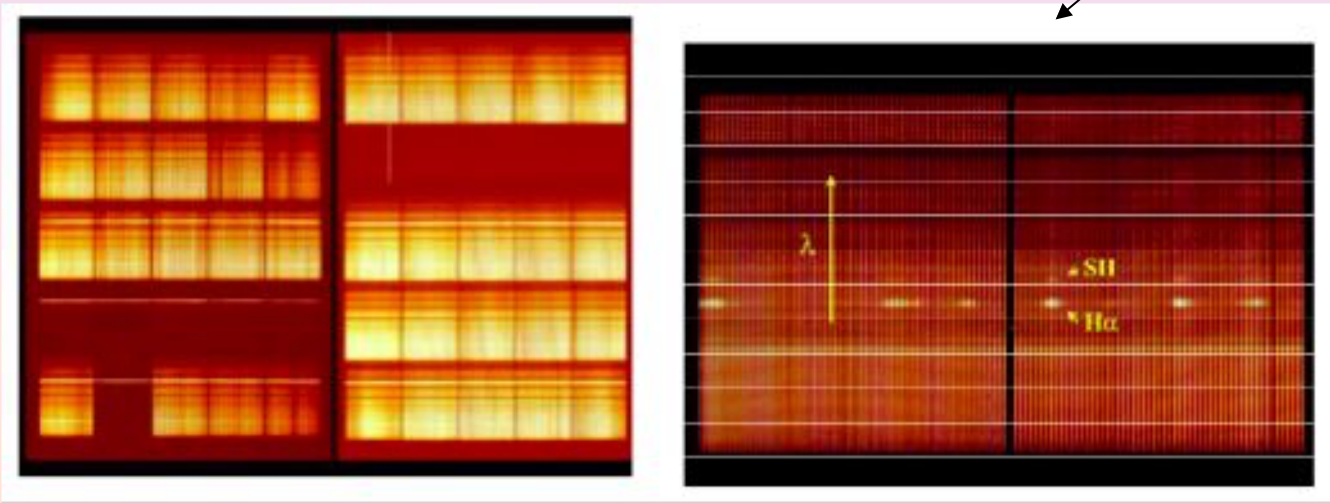
- pros and cons
  - Improves filling factor to near unity
  - Allows for control of input and output *f*-ratio
    - Effective coupling of slow telescope *f*-ratio to fiber input
    - Effective coupling of fiber output to spectrograph
  - May introduce scattered light (depends on lenslet)
  - Lower throughput (reflection, scattering, misalignment)
  - For science premium on truly integral field, above two factors don't out-weight filling factor improvements.
  - More subtle effect: *output lenslets?*
    - far-field vs near-field: where to control systematics:
    - Is your spectrograph seeing-limited or aberration limited?

Grating-dispersed spectrographs
fibers + lenslet instruments

- VIMOS, VLT 8m
  - 6400 or 1600 elements
  - 0.33 or 0.67 arcsec sampling
    - 13x13 arcsec up to 54x54 arcsec FoV
  - 360-1000 nm range
  - R = 200-2500

These are just two of 4 channels of 1600 fibers

These are just two of 20 groups of 80 fibers.



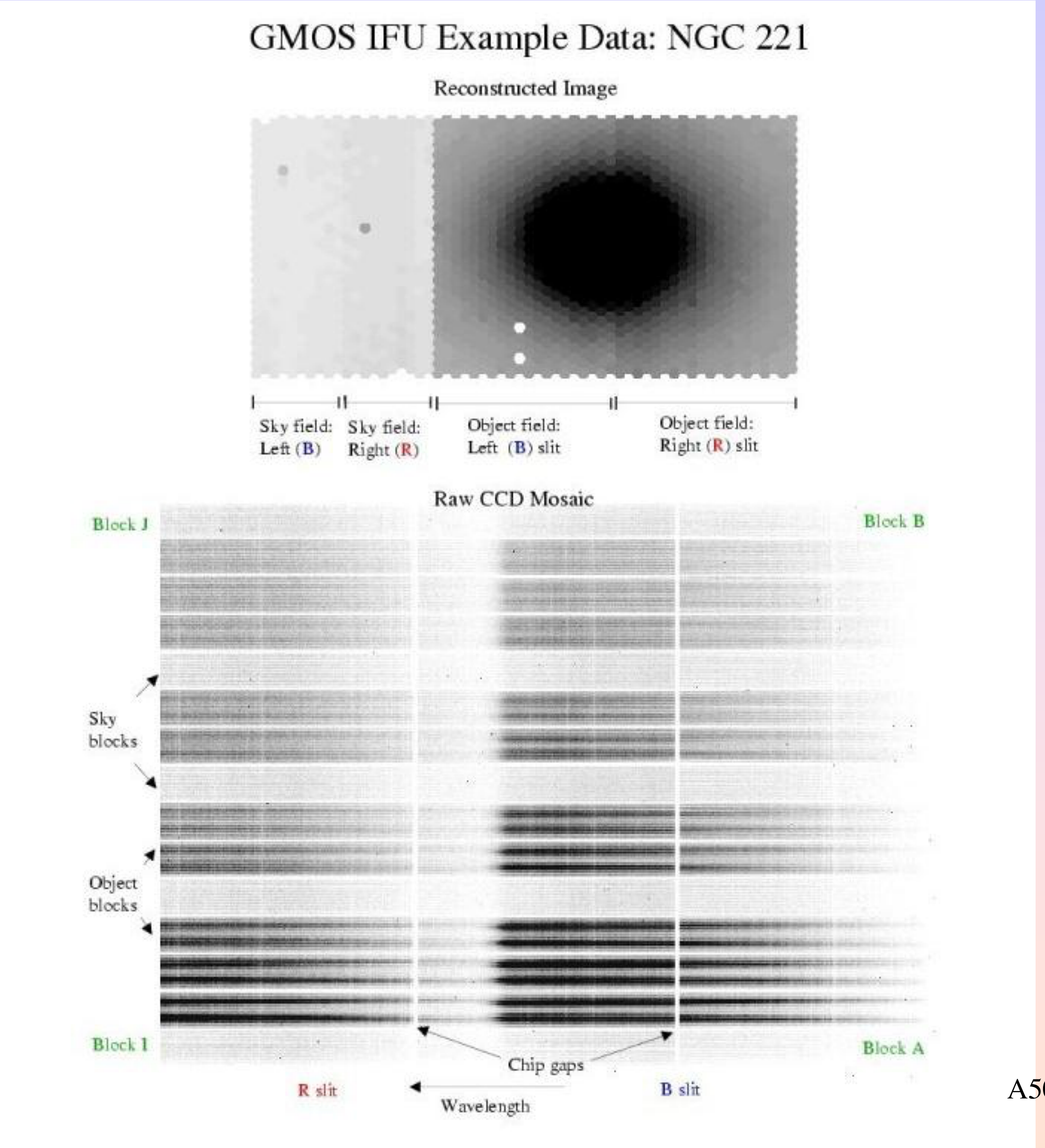
Le Fevre et al. 2003

commissioning data

Grating-dispersed spectrographs
fibers + lenslet instruments

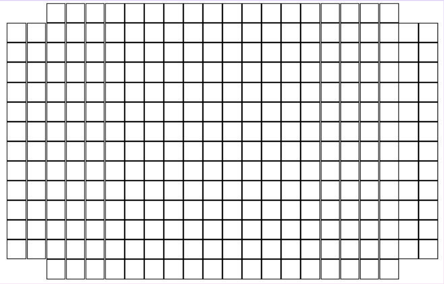
- GMOS, Gemini 8m
  - ➢ 1500 lenslets: 1000 object + 500 sky bundles
  - ➢ 0.2 arcsec per lenslet (7x5 arcsec + 5x3.5 arcsec FoV)
  - ➢ Two-slit and one-slit modes
  - ➢ R = 800-3500 (claim 10,000 achievable)
  - ➢ 400-1100 nm range

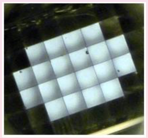
Allington-Smith et al.



Grating-dispersed spectrographs
fibers + lenslet instruments

- FLAMES/GIRAFFE: ARGUS, VLT 8m
  - 22x14 rectangular array
    - 0.52 arcsec per lens: 11.5x7.3 arcsec FoV
    - 0.3 arcsec per lens: 6.6x4.2 arcsec FoV
  - 15 20-element IFUs
    - 0.52 arcsec sampling (2x3 arcsec FoV)
  - R = 11,000-39,000 !
  - 370-950 nm range





(more about this instrument later)

Grating-dispersed spectrographs
fibers+lenslet instruments - summary list

- Existing optical instruments
  - PMAS, Calar Alto 3.5m
  - Spiral B, AAT 3.9m
  - MPFS, SAO 6m
  - IMACS-IFU, Magellan 6.5m
  - GMOS, Gemini 8m
  - VIMOS, VLT 8m
  - FLAMES/GIRAFFE ARGUS/IFU, VLT 8m
- Future optical instruments
- Existing infrared instruments
  - COHSI, UKIRT 3.8m (defunct?)
  - SMIRFS, UKIRT 3.8m (defunct?)
  - CIRPASS, Gemini 8m (defunct?)
- Future NIR instruments

No future?
Hard to believe.

Grating-dispersed spectrographs

fibers+lenslet instruments - summary list

Table 2. Fiber+Lenslet Integral Field Instruments

Instrument	Coupling Method	Telescope	D_T (m)	Ω (arcsec^2)	$d\Omega$ (arcsec^2)	N_θ	$\Delta\lambda/\lambda$	R	N_R	ϵ
Existing Optical Instruments										
PMAS	lenslet+fiber	Calar Alto	3.5	64.	0.5	256	0.11	9400.	1000	0.15
		Calar Alto	3.5	64.	0.5	256	0.52	1930.	1000	0.15
		Calar Alto	3.5	144.	0.75	256	0.11	9400.	1000	0.15
		Calar Alto	3.5	144.	0.75	256	0.52	1930.	1000	0.15
		Calar Alto	3.5	256.	1.0	256	0.11	9400.	1000	0.15
		Calar Alto	3.5	256.	1.0	256	0.52	1930.	1000	0.15
SPIRAL-B	lenslet+fiber	AAT	3.9	251.	0.49	512	0.29	1700.	495	···
		AAT	3.9	251.	0.49	512	0.07	7500.	495	···
MPFS	lenslet+fiber	SAO	6.0	256.	1.0	256	0.12	8800.	1024	0.045
		SAO	6.0	64.	0.25	256	0.47	2200.	1024	0.045
IMACS-IFU	lenslet+fiber	Magellan	6.5	···	···	···	···	···	···	···
GMOS	lenslet+fiber	Gemini	8.0	49.6	0.04	1500	0.21	3450.	730.	···
		Gemini	8.0	49.6	0.04	1500	0.32	2300.	730.	···
		Gemini	8.0	49.6	0.04	1500	0.82	890.	730.	···
		Gemini	8.0	24.8	0.04	750	0.42	3450.	1460.	···
		Gemini	8.0	49.6	0.04	1500	0.64	2300.	1460.	···
		Gemini	8.0	49.6	0.04	1500	1.00	890.	1460.	···
VIMOS	lenslet+fiber	VLT	8.0	2916.	0.45	6400	0.6	250.	150	···
		VLT	8.0	698.	0.11	6400	0.6	250.	150	···
		VLT	8.0	729.	0.45	1600	0.2	2500.	500	···
		VLT	8.0	174.5	0.11	1600	0.2	2500.	500	···
ARGUS/IFU	lenslet+fiber	VLT	8.0	83.9	0.27	315	0.105	11000.	1155	···
		VLT	8.0	83.9	0.27	315	0.042	39000.	1625	···
ARGUS	lenslet+fiber	VLT	8.0	27.7	0.09	315	0.105	11000.	1155	···
		VLT	8.0	27.7	0.09	315	0.042	39000.	1625	···
Future Optical Instruments										
Existing Near-Infrared Instruments										
COHSI	lenslet+fiber	UKIRT	3.8	···	···	100	0.26	500.	128	···
SMIRFS	lenslet+fiber	UKIRT	3.8	24.2	0.34	72	0.023	5500.	128	···
CIRPASS	lenslet+fiber	Gemini	8.0	54.5	0.13	490	0.41	2500.	1024	···
		Gemini	8.0	54.5	0.13	490	0.085	12000.	1024	···
		Gemini	8.0	27.0	0.06	490	0.41	2500.	1024	···
		Gemini	8.0	27.0	0.06	490	0.085	12000.	1024	···
Future Near-Infrared Instruments										

Grating-dispersed spectrographs
image slicer feeds

- Concept
  - S1: Slicer mirrors at telescope focal plane divide it into strips; mirrors have power to place telescope pupil on next element.
  - S2: Pupil mirrors (one per slice) reformat the slices into a pseudo-slit, where they form an image of the sky.
  - S3: Field lens control location of pupil stop in spectrograph.
- All-mirror and catadioptric designs exist.

The so called "advanced image slicer" →

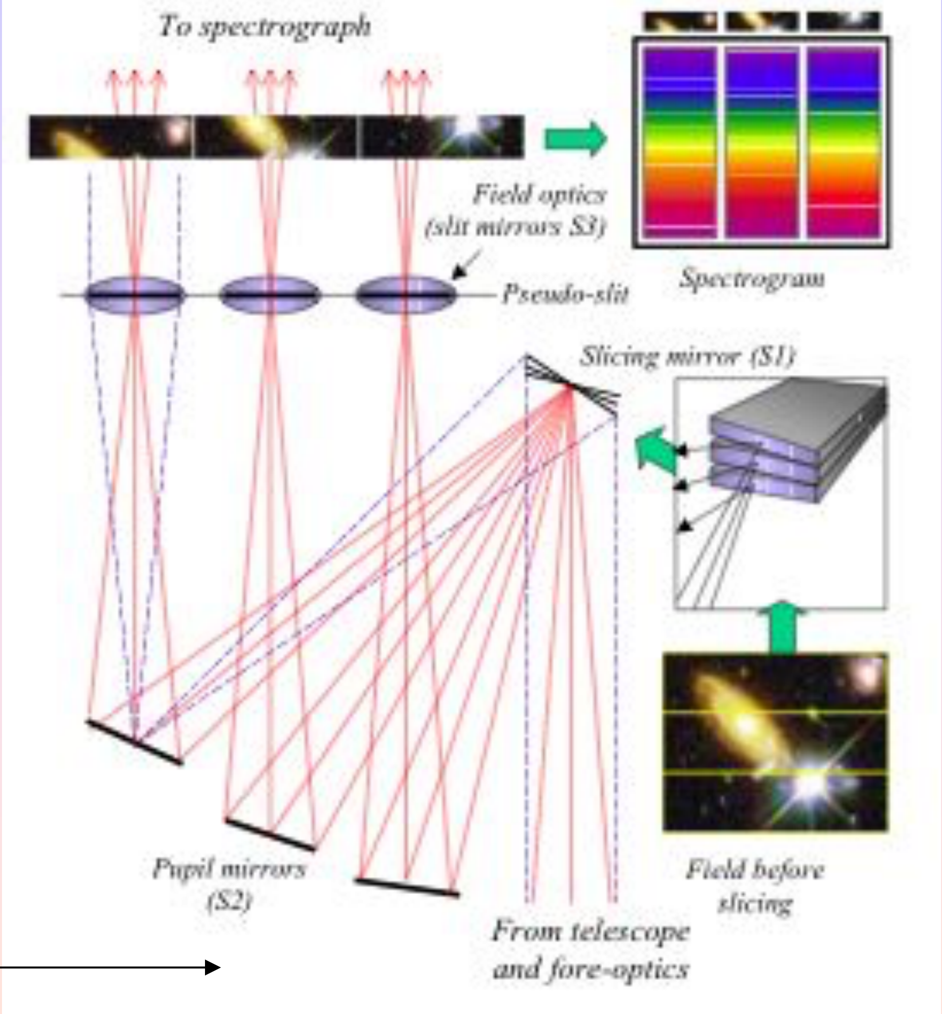


Figure 4. Image slicer principle (courtesy J. Allington-Smith, Durham U.)

Grating-dispersed spectrographs

image slicer feeds

- All mirror design for FISICA
  - ➢ Eikenberry et al. '04
    - o For FLAMINGOS
    - o 16x33 arcsec FoV on KPNO 4m
    - o 6x12 arcsec FoV on GTC
    - o JHK bands
    - o R = 1300

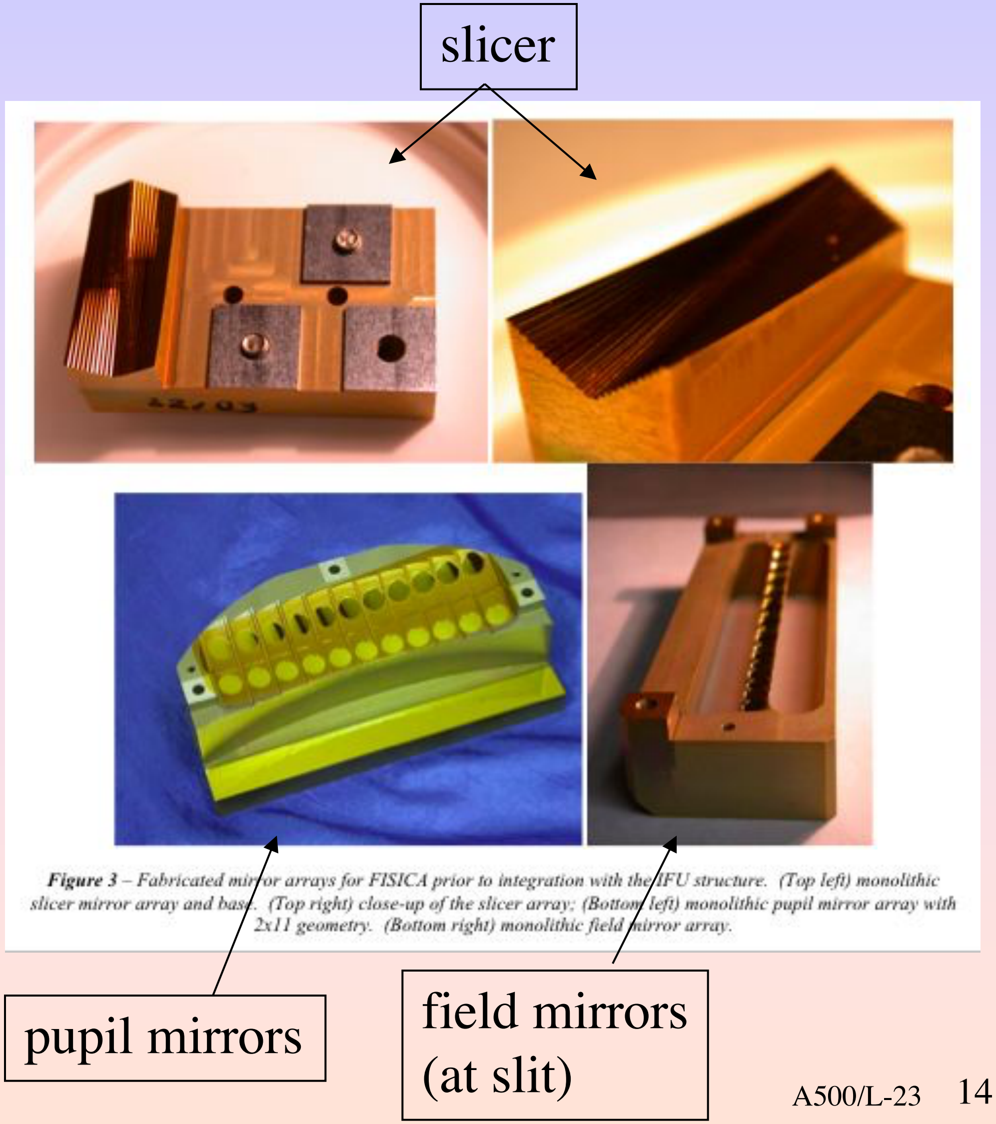


Figure 3 *– Fabricated mirror arrays for FISICA prior to integration with the IFU structure. (Top left) monolithic slicer mirror array and base. (Top right) close-up of the slicer array; (Bottom left) monolithic pupil mirror array with 2x11 geometry. (Bottom right) monolithic field mirror array.*

Grating-dispersed spectrographs
image slicer feeds

- Catadioptric Image Slicer (CIS) for MUSE

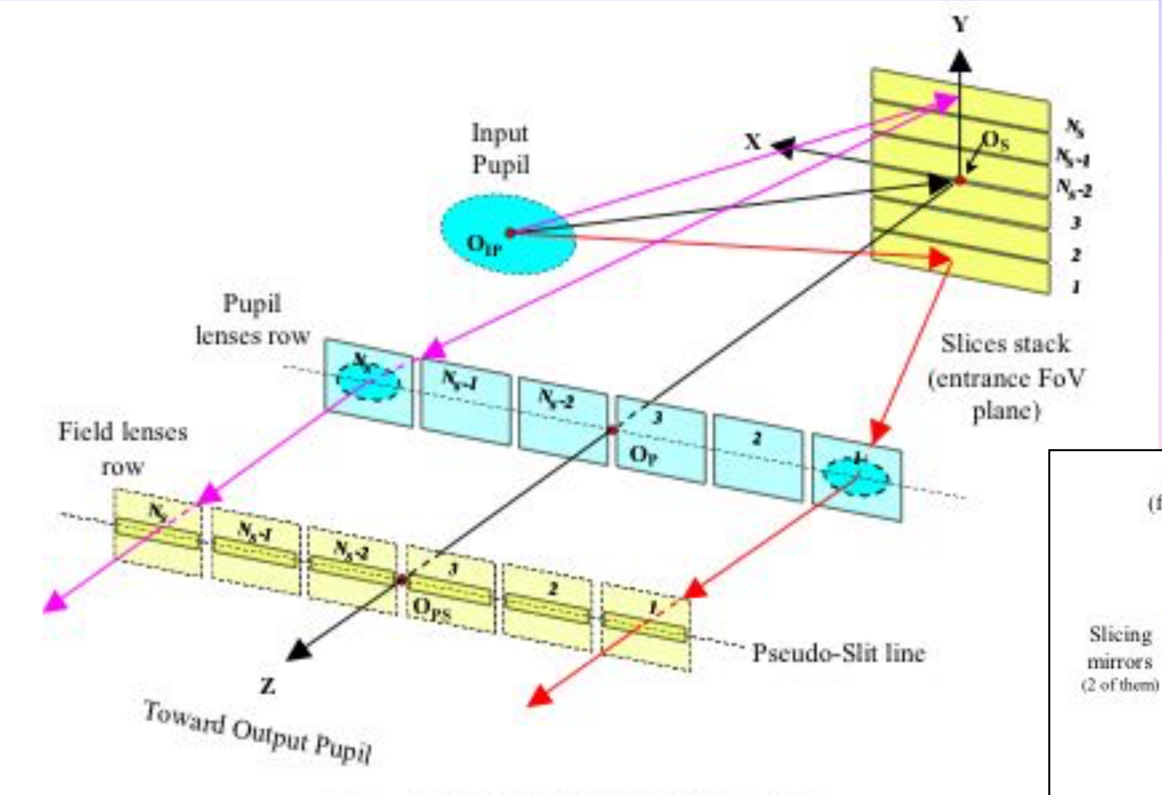


Figure 4-1: Principle of a Catadioptric Image Slicer

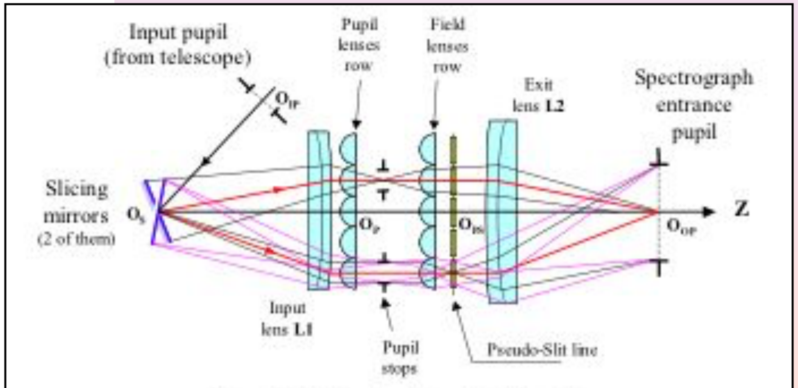


Figure 4-3: Final optical principle of the CIS

Henault et al. '03

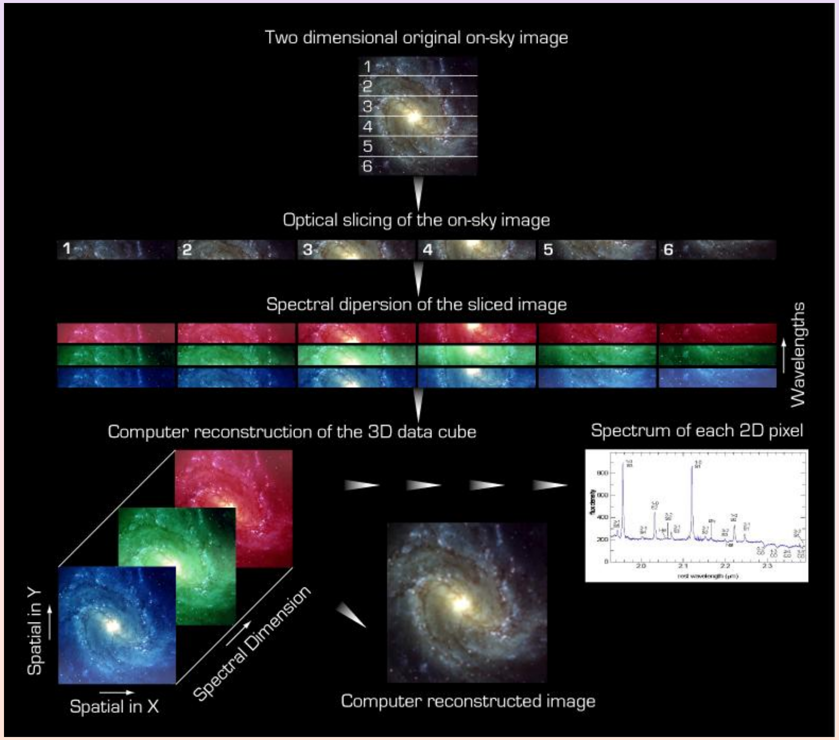
Grating-dispersed spectrographs
image slicer feeds

- Pros and cons
  - Image slicers are the only IFU mode to preserve all spatial information.
    - All other modes destroy spatial information within sampling element.
  - Most compact at reformating the focal plane onto detector.
  - Can be used in cryogenic systems and at long wavelengths where fibers don't transmit.
    - Lenslet arrays also accomplish this, e.g., OSIRIS (Keck)
  - Issues of scattered-light with diamond-turned optics mean optical slicers must be made differently.

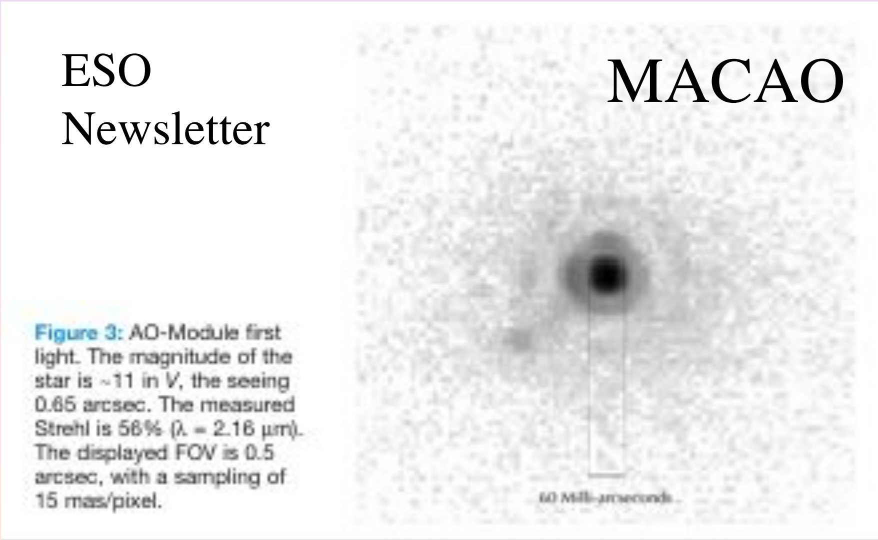
Grating-dispersed spectrographs
image slicer instruments

- SINFONI: SPIFFI + MACAO, VLT 8m
  - ➢ The power of near-infrared AO coupled to an image-slicing spectrograph
    - o 32x32 element imaging field sliced into a 1024 element long-slit
    - o Field coverage of 8x8 to 0.8x0.8 arcsec
    - o JHK coverage at R = 2000-4000

Bonnet et al. ‘04, Iserlohe et al. ‘04



SPIFFI



Grating-dispersed spectrographs

image slicer instruments

- SPIFFI slicer: pupil mirrors are flat

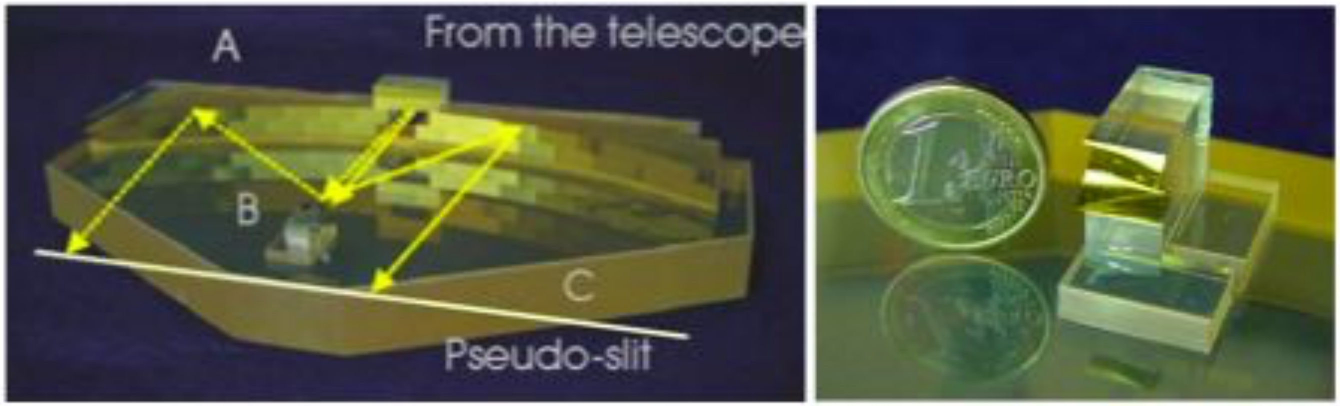


Figure 2. The image slicer: The image slicing device used in SPIFFI consists of plane mirrors. The two dimensional image (still containing light of all wavelengths of the used observing band) is sliced into thin stripes, called slitlets, which are then lined up end to end to create a pseudo long slit which is fed into the spectrometer. The small image slicer (B in left Figure) is a stack of 32 plane mirrors with a thickness of 0.3mm each. The small image slicer (right Figure) cuts the image into stripes and reflects them onto the big image slicer (A) which consists of 32 plane mirrors in two layers. The mirrors are tipped and tilted in a way that the telecentric entrance pupil is preserved. Both image slicer components are mounted to a baseplate(C). All parts are of Zerodur and are optically contacted (without using any glue).

Iserlohe et al. '04

Grating-dispersed spectrographs
image slicer feeds

- Catadioptric Image Slicer for SINFONI/SPIFFO

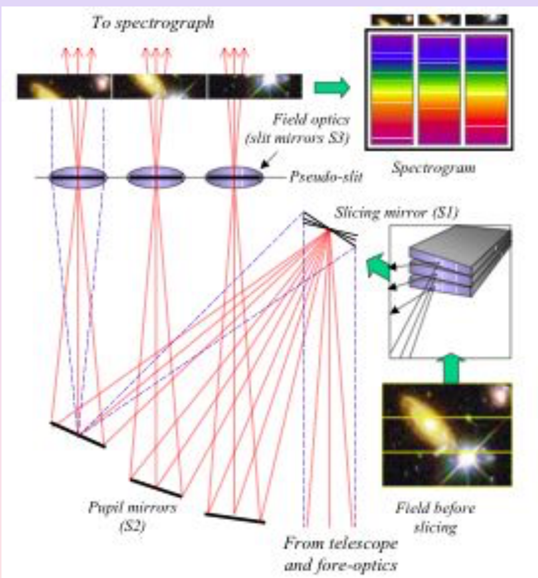
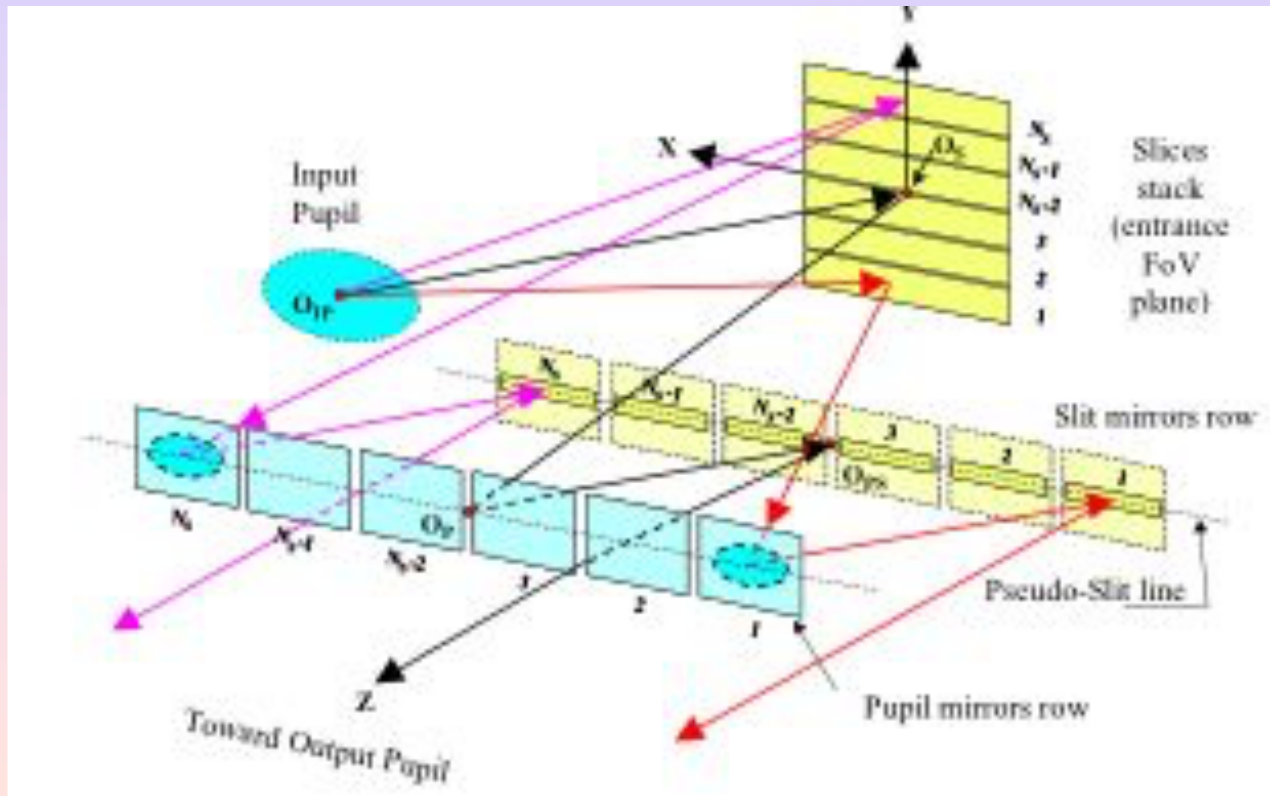


Figure 4. Image slicer principle (courtesy J. Allington-Smith, Durham U.)



Henault et al. '03

Grating-dispersed spectrographs
image slicer instruments

Note shifts in adjacent slices: different field-angles from two slicer layers

- SPIFFI data and format

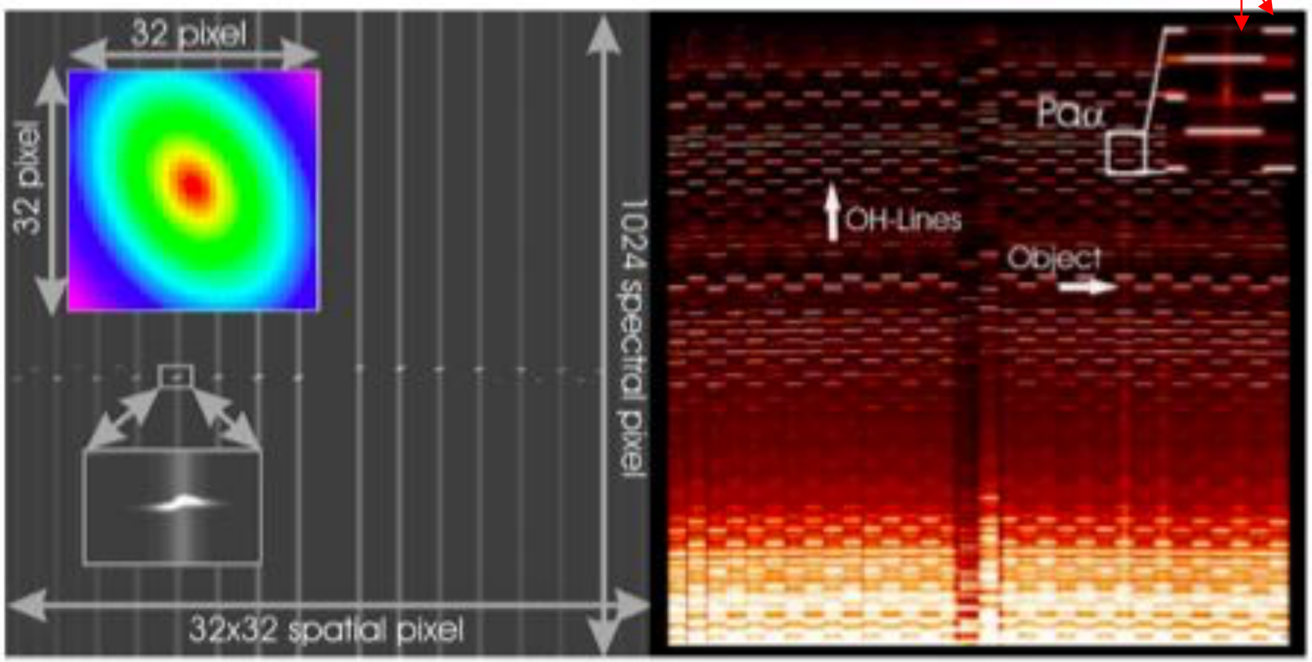
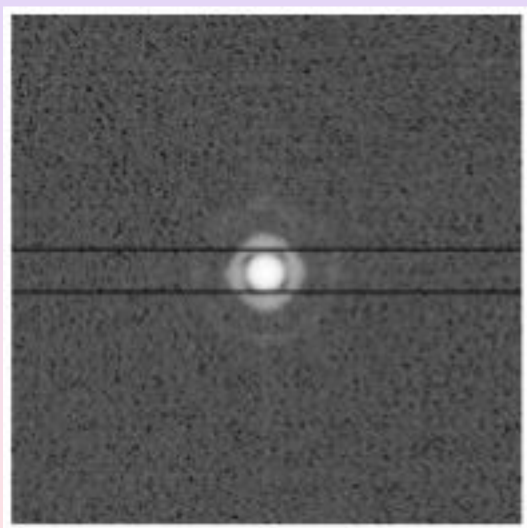


Figure 3. Left: Simulated image of a rotating source with continuum and one emission line. Each slitlet is 32 columns wide on the detector and with 32 slitlets one obtains 1024 spatial pixels with each spectrum covering exactly one detector column. The inset to the upper left shows the reconstructed image after having rearranged the slitlets. The inset below shows the velocity disperged emission line. Right: K-band detector raw frame of the ULIRG IRAS06206-6315. Together with bright night skylines (mainly OH), thermal atmospheric background, continuum and line emission of the source can be identified (long wavelengths at the bottom). The inset shows Pa_α line emission.

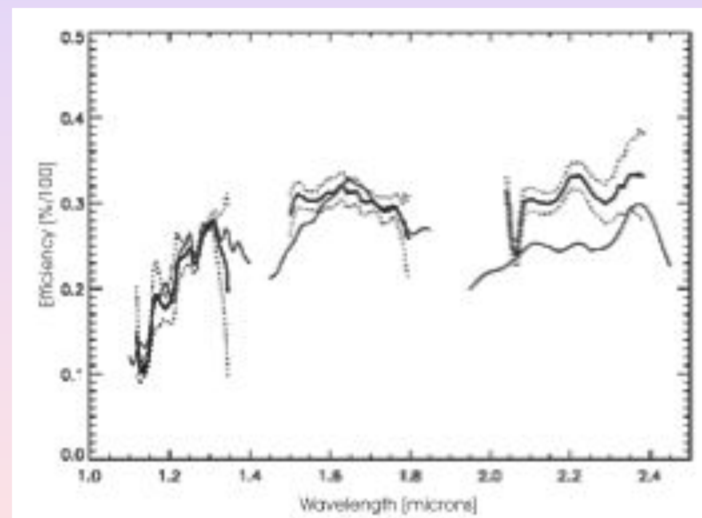
Iserlohe et al. '04

Grating-dispersed spectrographs
image slicer instruments

- SINFONI efficiencies:



Slicer mirror width

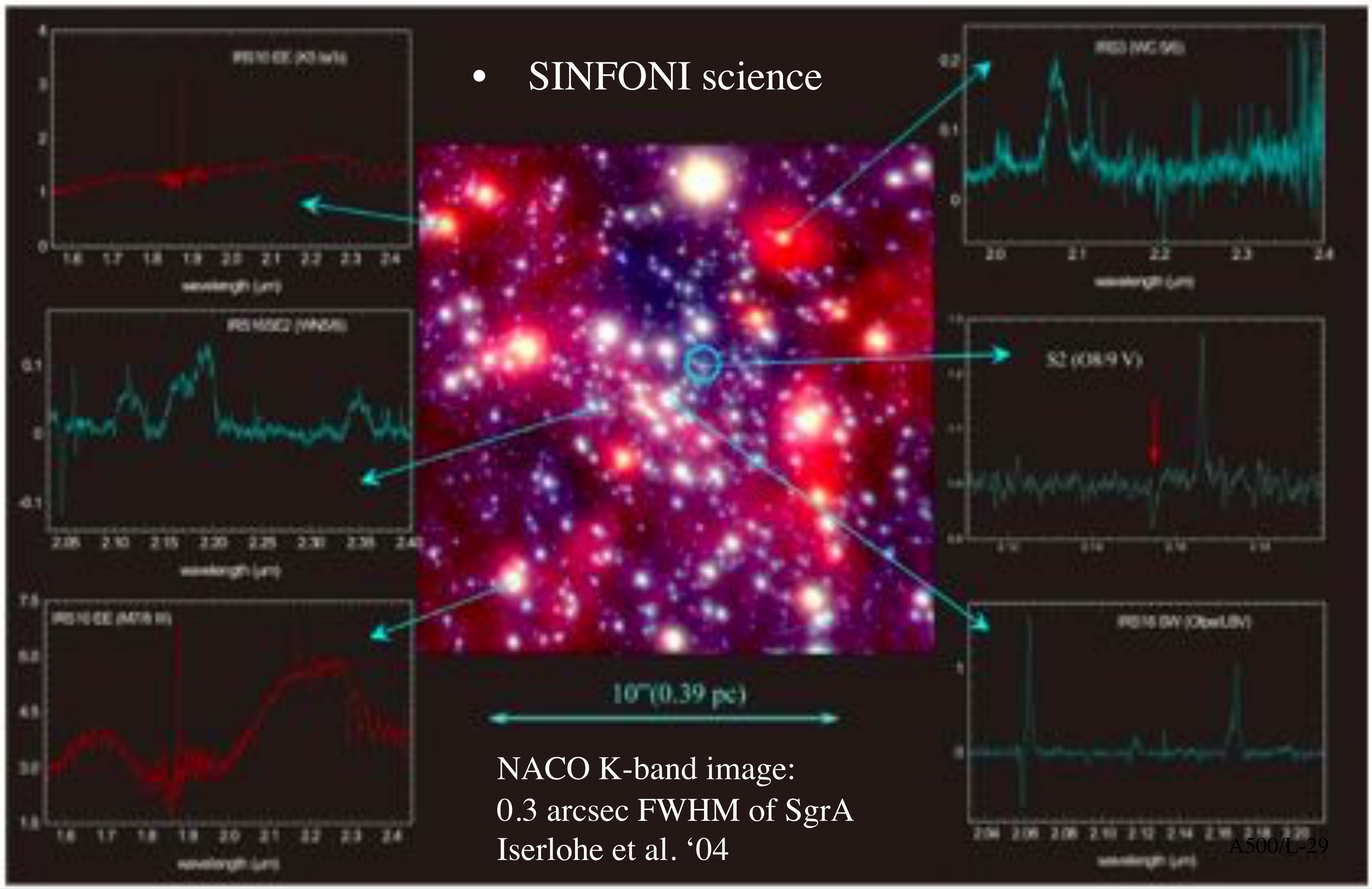


Total system throughput 20-30%

Iserlohe et al. ‘04

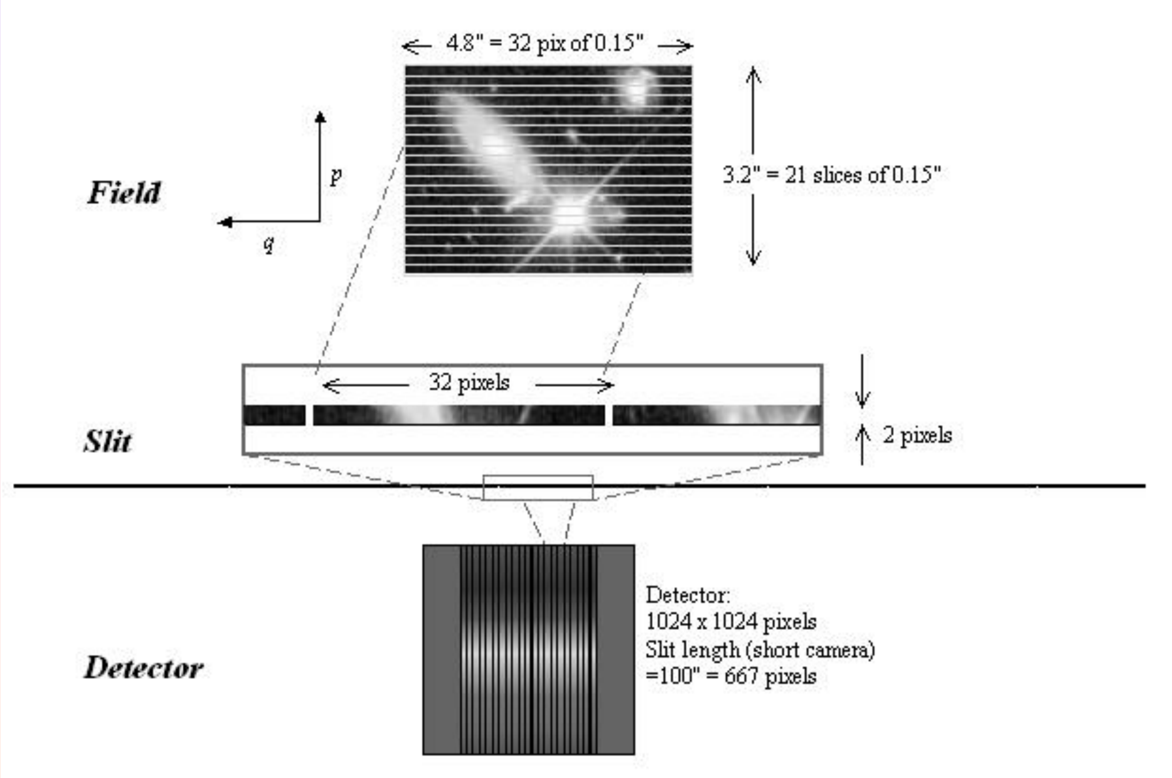
Grating-dispersed spectrographs

image slicer instruments



Grating-dispersed spectrographs
image slicer instruments

- GNIRS, Gemini-S, 8m
  - 4.8x3.2 arcsec FoV (0.15 arcsec slices)
  - R = 1700 and 5900
  - 0.9-5.5 μm range (2.5-5.5 μm not commissioned)



Grating-dispersed spectrographs
image slicer instruments

- GNIRS, Gemini-S, 8m
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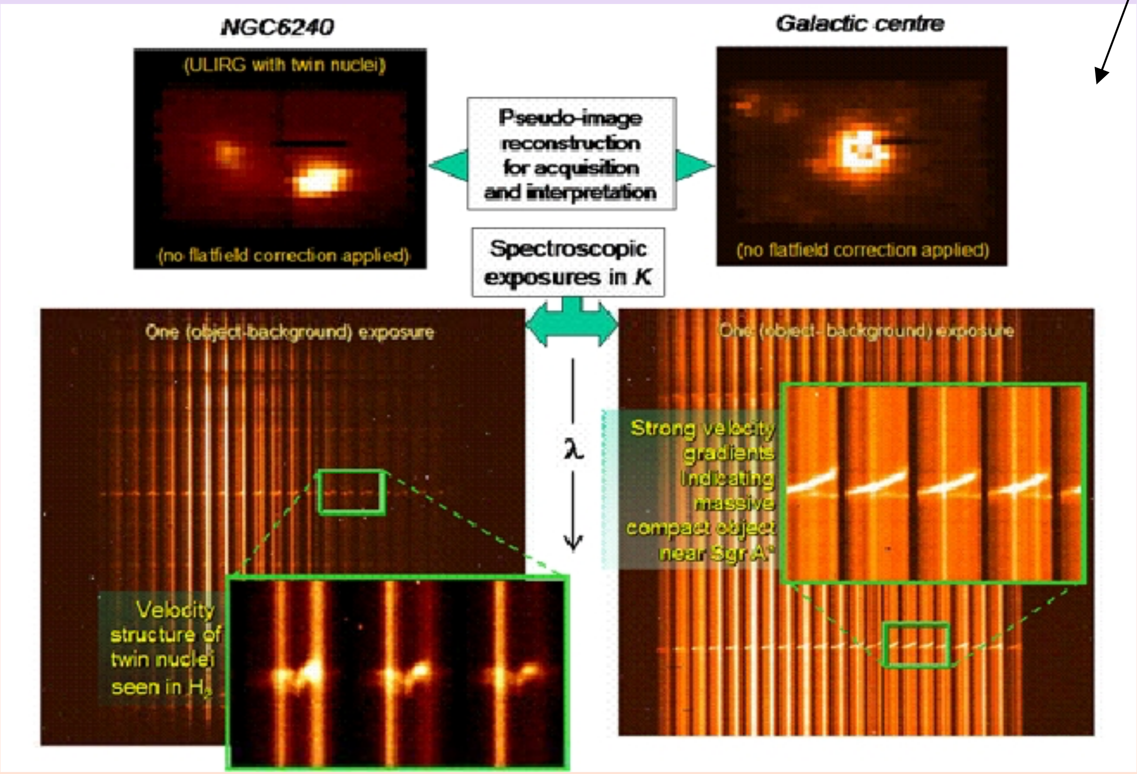
Allington-Smith et al. ‘04
Commissioning data



Grating-dispersed spectrographs
image slicer instruments - summary list

- Optical instruments
  - WiFeS, ANU 2.3m
  - SWIFT, Palomar 5m (see: Goodsall poster)
  - GISMO, Magellan 6.5m
  - **MUSE, VLT 8m**
- Existing NIR instruments
  - UIST, UKIRT 3.8m
  - MPE-3D, AAT 3.9m (defunct)
  - PIFS, Palomar 5m
  - <u>GNIRS, Gemini 8m</u>
  - SPIFFI, VLT 8m
  - <u>NIFS, Gemini 8m</u>
- New NIR instruments
  - **KMOS, VLT 8m**
  - FISICA, GTC 10.4m
- New NIR instruments in space
  - <u>**NIRSpec, JWST 6.5m**</u>
  - <u>**MIRI, JWST 6.5m**</u>
- Future NIR instruments
  - NIRMOS, GMT 25.4m

Note extensive NIR instrumentation; relatively few in optical because surface smoothness requirements are higher.

Grating-dispersed spectrographs
image slicer instruments - summary list

Table.3 Slicer Integral Field Instruments

Instrument	Coupling Method	Telescope	D_T (m)	Ω (arcsec^2)	$d\Omega$ (arcsec^2)	N_θ	$\Delta\lambda/\lambda$	R	N_R	ϵ
				Existing Optical Instruments						
ESI	slicer	Keck	10.0	···	···	···	···	···	···	···
				Future Optical Instruments						
WiFeS	slicer	ANU	2.3	775.	1.	775	1.03	3000.	3090	···
		ANU	2.3	775.	1.	775	0.44	7000.	3090	···
IMACS/GISMO	slicer	Magellan	6.5	···	···	···	···	···	···	···
MUSE	advanced-slicer	VLT	8.0	3600	0.04	9e4	0.67	3000.	2000	0.24
				Existing Near-Infrared Instruments						
PIFS	slicer	Palomar	5.0	51.8	0.45	115	0.23	550.	128	0.22
		Palomar	5.0	51.8	0.45	115	0.10	1300.	128	0.22
GNIRS	advanced-slicer	Gemini	8.0	15.4	0.023	684	0.301	1700.	512	···
		Gemini	8.0	15.4	0.023	684	0.087	5900.	512	···
SPIFI	slicer	VLT	8.0	0.54	0.006	1024	0.34	3000.	1024	0.3
		VLT	8.0	10.2	0.001	1024	0.34	3000.	1024	0.3
		VLT	8.0	64.0	0.06	1024	0.34	3000.	1024	0.3
NIFS	advanced-slicer	Gemini	8.0	9.0	0.01	900	0.19	5300.	1007	···
				Future Near-Infrared Instruments						
KMOS	advanced-slicer	VLT	8.0	188.0	0.04	4204	0.28	3600.	1000	···
FISICA/FLMINGOS	advanced-slicer	GTC	10.4	72.0	0.53	136	0.79	1300.	1024	···
				Future Optical–Near-Infrared Space-Based Instruments						
NIRSpec	advanced-slicer	JWST	6.5	···	···	···	···	···	···	···
MIRI	advanced-slicer	JWST	6.5	···	···	···	···	···	···	···
SNAP	advanced-slicer	SNAP	2	···	···	···	···	···	···	···

Grating-dispersed spectrographs
lenslet feed

- Concept: pupil imaging spectroscopy using lenslets. No fibers to reformat telescope focal plane into long-slit.

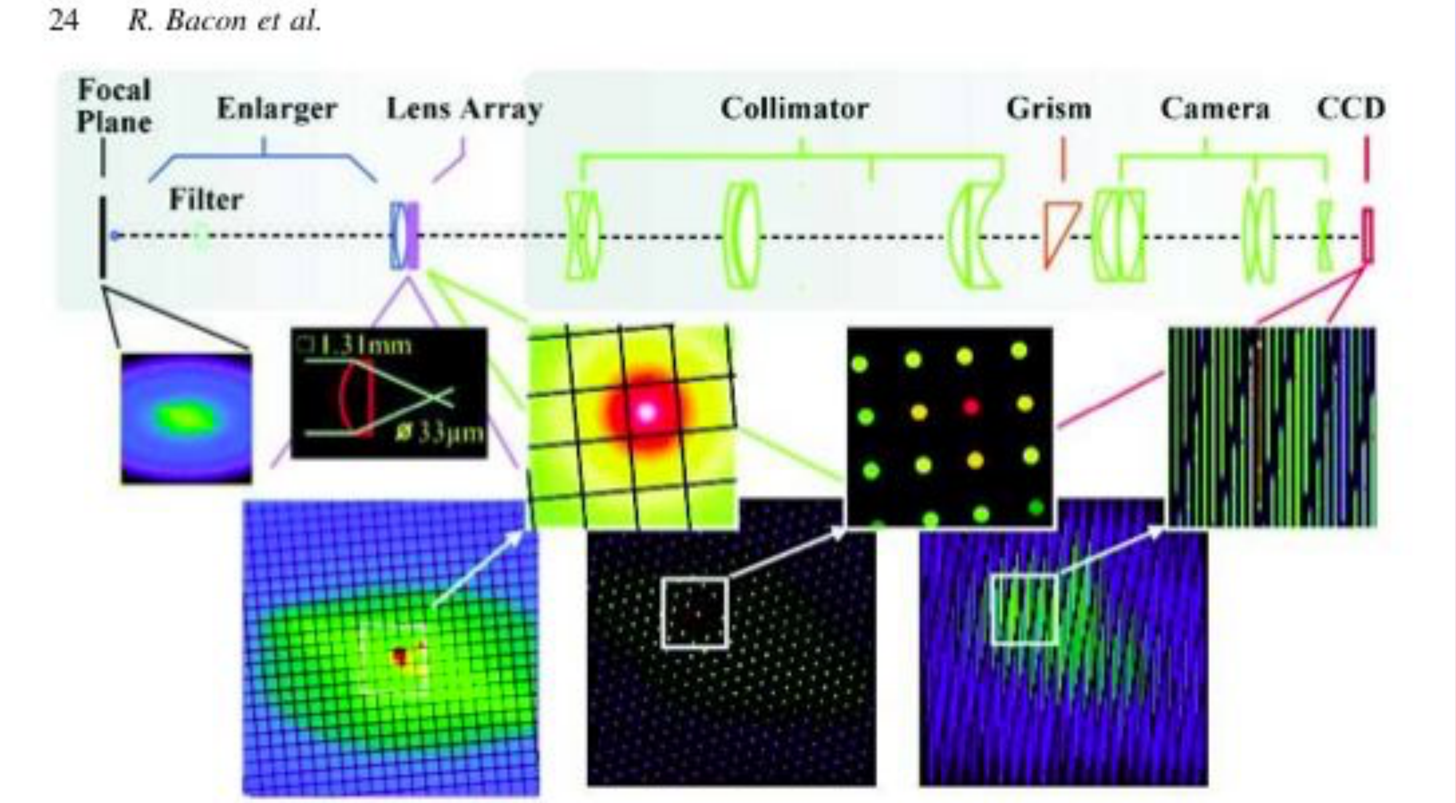
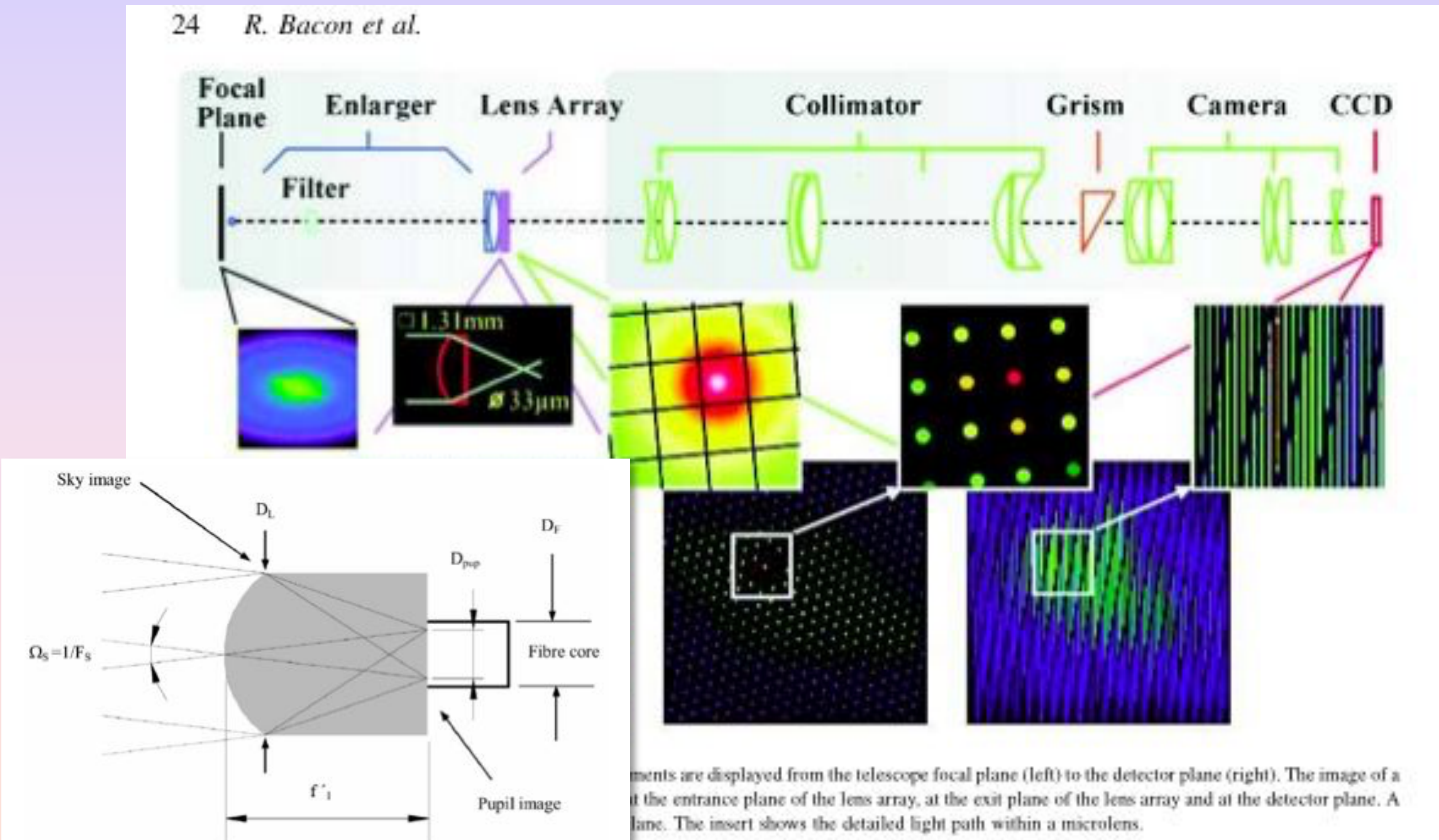


Figure 1. Optical layout of SAURON. The main optical elements are displayed from the telescope focal plane (left) to the detector plane (right). The image of a galaxy is shown successively at the telescope focal plane, at the entrance plane of the lens array, at the exit plane of the lens array and at the detector plane. A zoom-in of a small area of the galaxy is shown for each plane. The insert shows the detailed light path within a microlens.

SAURON: Bacon et al. '01 "1577 eyes see all"

Grating-dispersed spectrographs
lenslet feed

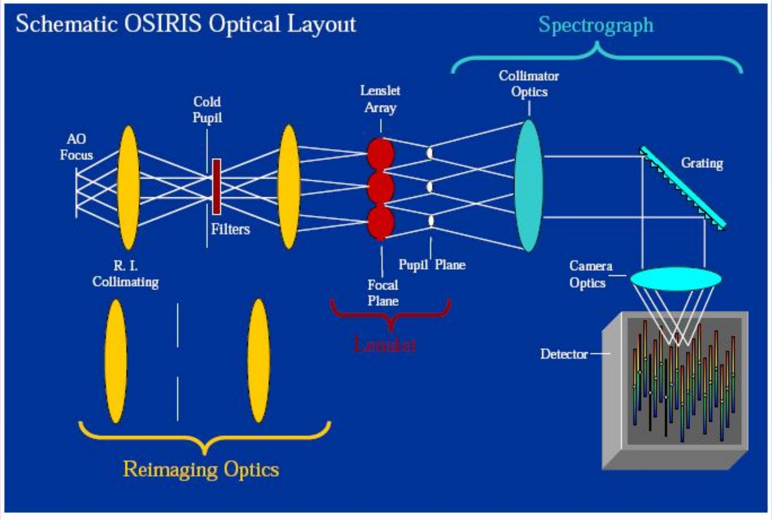
- Concept: pupil imaging spectroscopy using lenslets. No fibers to reformat telescope focal plane into long-slit.



SAURON: Bacon et al. ‘01 “1577 eyes see all”

Grating-dispersed spectrographs
lenslet feed: pros and cons

- Excellent fill factor
- Low dispersion due to grism limitations
- Grating implementation now exists in the near-infrared:
- VPH grisms and gratings with articulated camera would yield higher resolution
- implications of formatting for efficiency and scattered light
  - good data packing on detector, but limited spectral coverage/resolution
  - scattered light from lenslets may be significant



OSIRIS (Keck 10m)

Larkin et al.'03

Grating-dispersed spectrographs

lenslets: formatting

28 *R. Bacon et al.*

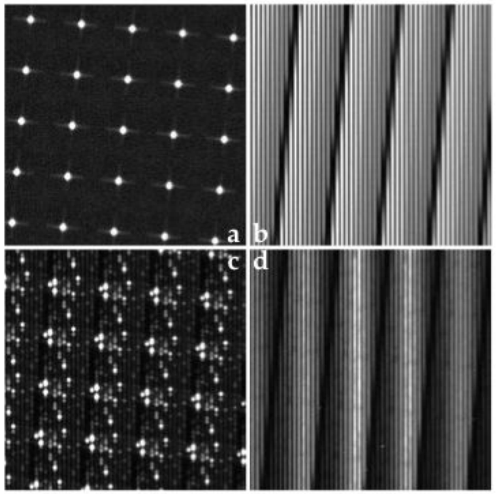


Figure 5. Examples of SAURON exposures taken during commissioning. Each panel shows only a small part of the entire CCD frame so that details can be seen. (a) Micropupil image taken with the grism out. Each cross is the diffraction pattern from one square lenslet. (b) Continuum image using the tungsten lamp. (c) Neon arc. (d) Central part of NGC 3377. This figure is available in colour in the electronic version of the article on *Synergy*.

- Spectral extraction critical to minimize crosstalk

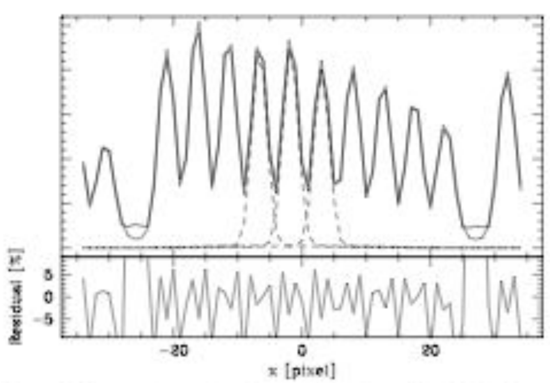


Figure 8. Upper panel: example of a cross-dispersion profile derived from a continuum exposure (thick solid line) and its fit by the model described in text (thin solid line). To illustrate the level of overlap, only the central three fitted profiles are shown (dashed lines). Lower panel: relative residual of the fit (in per cent).

SAURON

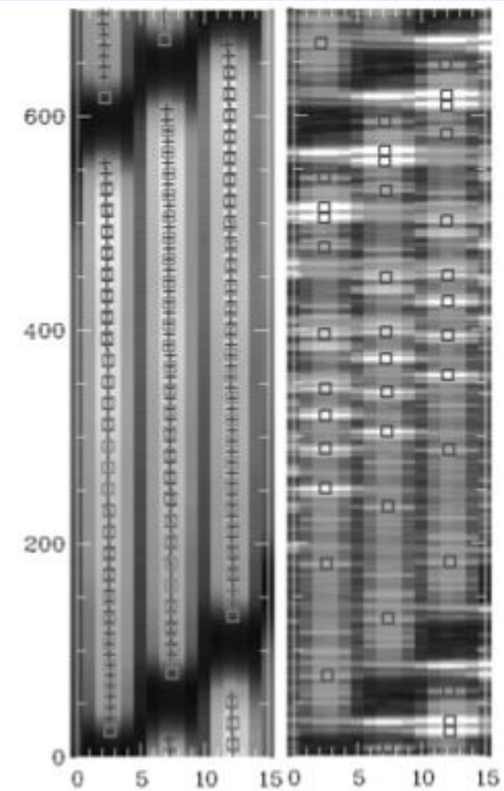
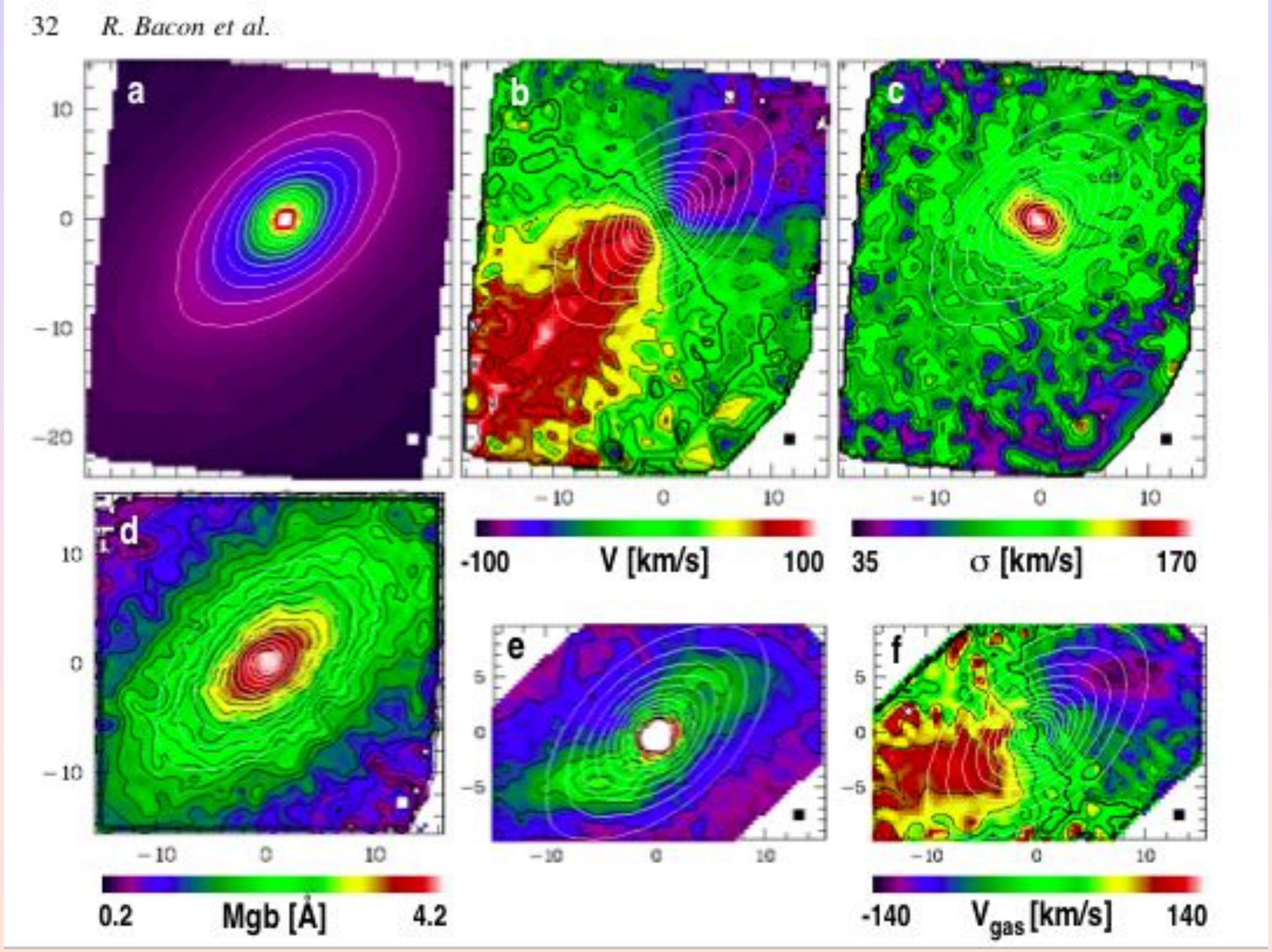


Figure 9. Example of mask-fitting results. Left panel: a small part of a continuum exposure with the derived locations of the spectra (crosses) and the corresponding fitted values (squares). Right panel: the corresponding arc (neon) exposure and fitted emission lines (squares). The images have been expanded along the cross-dispersion axis for clarity.

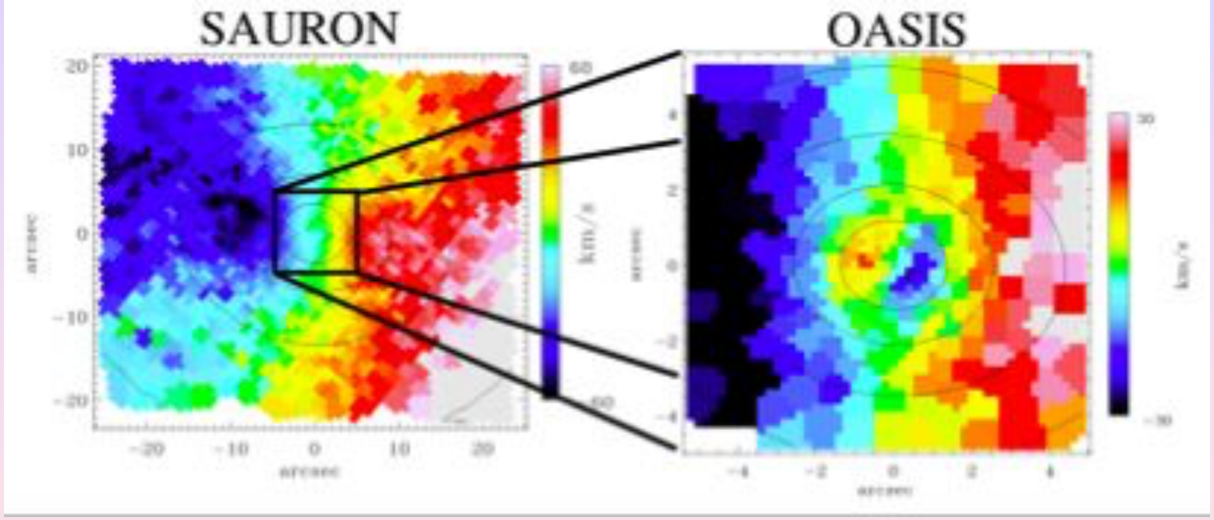
Grating-dispersed spectrographs
lenslets science

- impressive science results from SAURON



Grating-dispersed spectrographs
lenslets science

- The power of AO coupling to truly integral field: OASIS on WHT



McDermid et al. ‘04

Grating-dispersed spectrographs
lenslet instruments - summary list

- Existing optical instruments
  - OASIS, WHT 4.2m (formerly on CFHT 3.5m)
  - SAURON, WHT 4.2m
- Future optical instruments
- Existing NIR instruments
  - OSIRIS, Keck 10m
- Future NIR instruments
  - Planet Finder, VLT 8m

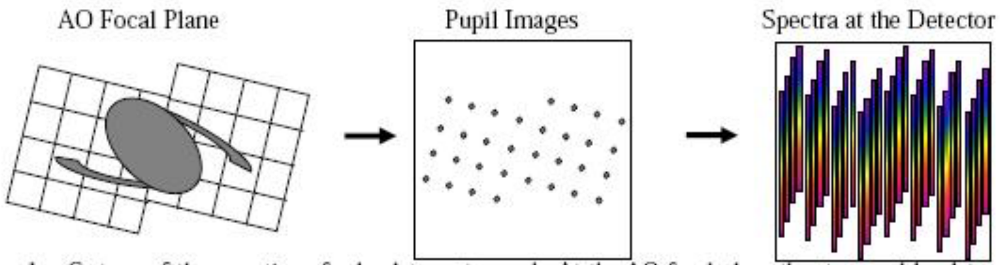


Figure 1 – Cartoon of the operation of a lenslet spectrograph. At the AO focal plane the staggered lenslet array samples the image and each lens produces a pupil image roughly one focal length behind the array. This plane is used to feed the collimator of a standard spectrograph. Since the pupils are well separated, the spectra can be interleaved between each other and can cover most of the detector.

Grating-dispersed spectrographs
lenslet instruments - summary list

Table 4. Lenslet Integral Field Instruments

Instrument	Coupling Method	Telescope	D_T (m)	Ω (arcsec^2)	$d\Omega$ (arcsec^2)	N_θ	$\Delta\lambda/\lambda$	R	N_R	ϵ
Existing Optical Instruments										
SAURON	lenslet	WHT	4.2	1353	0.88	1577	0.11	1213.	128	⋯
		WHT	4.2	99	0.07	1577	0.10	1475.	150	⋯
OASIS	lenslet	WHT	4.2	1.92	0.002	1100	0.50	1000.	400	⋯
		WHT	4.2	31.0	0.026	1100	0.50	1000.	400	⋯
		WHT	4.2	180.	0.17	1100	0.50	1000.	400	⋯
Future Optical Instruments										
Existing Near-Infrared Instruments										
OSIRIS	lenslet	Keck	10.4	1.2	0.02	3000	0.12	3400.	400	⋯
		Keck	10.4	30.	0.10	3000	0.12	3400.	400	⋯
		Keck	10.4	0.3	0.02	1019	0.47	3400.	1600	⋯
		Keck	10.4	7.5	0.10	1019	0.47	3400.	1600	⋯
Future Near-Infrared Instruments										